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SUPERNOVAE 2005eb and 2005ec

Further to *IAUC* 8598, N. Ponticello and W. Li report the LOSS/KAIT discovery of an apparent supernova (designated 2005eb) on unfiltered KAIT images taken on Sept. 7.33 (at mag 18.4) and 16.33 UT (mag 18.8) and located at $\alpha = 0^h54^m50^s.94$, $\delta = +29°14'46''.8$ (equinox 2000.0), which is $9''.0$ east and $2''.4$ south of the center of UGC 556. Nothing was visible at this location on a KAIT image taken on Aug. 27.32 (limiting mag 19.5).

Ponticello and Li also report the LOSS/KAIT discovery of another supernova (designated SN 2005ec) on unfiltered KAIT images taken on Sept. 16.54 (at mag 16.8) and 17.54 UT (mag 16.6) and located at $\alpha = 4^h54^m19^s.17$, $\delta = +1°38'27''.6$ (equinox 2000.0), which is $1''.3$ west and $2''.5$ north of the center of NGC 1690. Nothing was visible at this location on a KAIT image taken on Jan. 11.13 (limiting mag 19.0).

The "Nearby Supernova Factory" group (G. Aldering *et al.*; cf. *CBET* 230) reports that a spectrum (range 320–1000 nm) of SN 2005ec, obtained on Sept. 21.6 UT with the Supernova Integral Field Spectrograph on the University of Hawaii 2.2-m telescope, shows it to be a type-Ia supernova at an age of two weeks after maximum and at a redshift of $z = 0.03$, in agreement with the host redshift (NGC 1690; Ostriker *et al.* 1988, *A.J.* **96**, 1775). The spectrum is similar to that of SN 1994D at this phase, shifted to the host's redshift.

V838 MONOCEROTIS

M. Claussen, National Radio Astronomy Observatory (NRAO); K. Healy and S. Starrfield, Arizona State University; and H. E. Bond, Space Telescope Science Institute, report SiO maser observations of the unusual eruptive variable V838 Mon using the NRAO Very Large Array on Sept. 1: "SiO maser emission was first detected by Deguchi *et al.* (2005, `http://xxx.lanl.gov/abs/astro-ph/?0507363`) on 2005 Feb. 23, strengthening by Apr. 24. Relative to the April observations, we find that the peak flux in the v=2, J=1-0 SiO transition at 42.82 GHz has increased further by a factor of four to 4.98 ± 0.01 Jy. The peak flux in the v=1, J=1-0 transition at 43.12 GHz has also strengthened by a factor of three to 2.83 ± 0.01 Jy. The maser emission in both transitions is unresolved (angular size $< 0''.3$). Further monitoring with the VLA and the Very Long Baseline Array is planned, and monitoring is encouraged at other wavelengths, especially in the near- and mid-infrared."

2005 September 21 *Daniel W. E. Green*