

**Central Bureau for Astronomical Telegrams
INTERNATIONAL ASTRONOMICAL UNION**

Mailstop 18, Smithsonian Astrophysical Observatory, Cambridge, MA 02138, U.S.A.
IAUSUBS@CFA.HARVARD.EDU or FAX 617-495-7231 (subscriptions)
CBAT@CFA.HARVARD.EDU (science)
URL <http://cfa-www.harvard.edu/iau/cbat.html> ISSN 0081-0304
Phone 617-495-7440/7244/7444 (for emergency use only)

SUPERNOVAE 2005kd, 2005ke, 2005kf

Three apparent supernovae have been reported: 2005kd by T. Puckett and A. Pelloni (via the 0.35-m automated supernova patrol telescope; cf. *IAUC* 8629), and 2005ke and 2005kf by M. Baek, R. R. Prasad, and W. Li (via LOSS/KAIT; cf. *IAUC* 8628).

SN	2005 UT	α_{2000}	δ_{2000}	Mag.	Offset
2005kd	Nov. 12.22	4 ^h 03 ^m 16 ^s .88	+71°43'18".9	17.0	0".1 W, 5".0 N
2005ke	Nov. 13.33	3 35 04.35	-24 56 38.8	17.2	40" E, 40" S
2005kf	Nov. 11.53	7 47 26.51	+26 55 32.4	17.2	0".8 E, 0".6 S

Additional approximate magnitudes for 2005kd in PGC 14370: Sept. 18 and Nov. 9, [20.0; Nov. 13.02 UT, 16.7. Additional approximate magnitudes for 2005ke in NGC 1371: Oct. 24.32, [19.5; Nov. 14.35, 16.3. Additional approximate magnitudes for 2005kf, which is in a galaxy located 82" east and 16".5 south of the center of NGC 2449: Apr. 5.20, [19.5; Oct. 14.52, 17.0; 20.53, 16.5; 31.50, 16.7; Nov. 16.52, 17.5.

SN 2005kd has been found by an Ohio State University group (J. Eastman *et al.*, communicated by J. Prieto; cf. *CBET* 290) to be a young type-II_n supernova from a spectrogram (range 390–730 nm) taken on Nov. 13.3 UT with the MDM 2.4-m telescope (+ CCDs); the spectrum shows a blue continuum and strong hydrogen Balmer (H α , H β , H γ , and H δ) and He I lines in emission at central wavelengths fully consistent with the recession velocity of PGC 14370.

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D. Pray, Greene, RI; P. Pravec and P. Kušnirák, Ondřejov Observatory; and W. Cooney, J. Gross and D. Terrell, Sonoita Research Observatory, report that observations obtained during Nov. 1–13 show that the minor planet (2006) consists of a pair of bodies orbiting each other with a period of 19.15 ± 0.02 hr and producing mutual eclipse/occultation events 0.06-mag deep. There are present two rotational lightcurves with periods of 3.1183 ± 0.0002 and 6.656 ± 0.001 hr, with amplitudes of 0.08 and 0.06 mag, respectively, that persist during the mutual events. This suggests that either the system is binary with both bodies rotating non-synchronously with the orbital motion or there is a third body producing the additional period. A lower limit on the mean-diameter ratio of the eclipsing pair is 0.22. Further observations are needed to fully resolve the system.